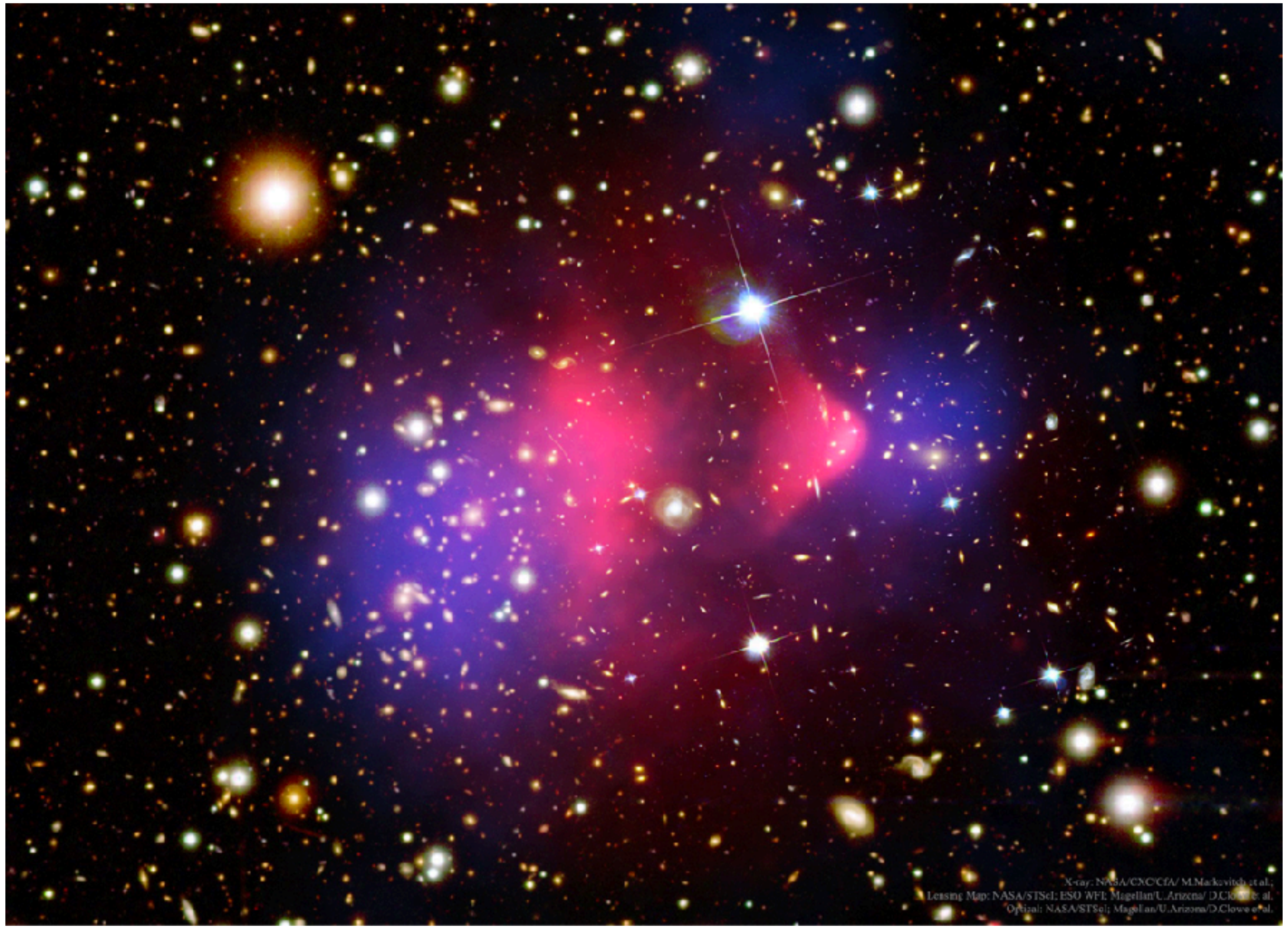
Dark Matter

ASTR/PHYS 4080: Intro to Cosmology

Week 7



How much visible matter is there?

Can only see matter that emits light.

Surveys tell us that in the local universe

$$(\text{out to } d \sim 0.1c/H_0)$$

the luminosity density in the V band is

$$\Psi_V = 1.1 \times 10^8 \ L_{\odot,V} \ \text{Mpc}^{-3}$$

where $L_{\odot,V} \approx 0.12 L_\odot \approx 4.6 \times 10^{25} \ \text{W}$

But we want their mass, which we can infer if we know the typical *mass-to-light ratio*

Because astronomers do things in relative terms, we compare to the Sun:

$$\langle M/L_V \rangle = 1 \ M_\odot/L_{\odot,V}$$

But of course, different stars have different M/L values:

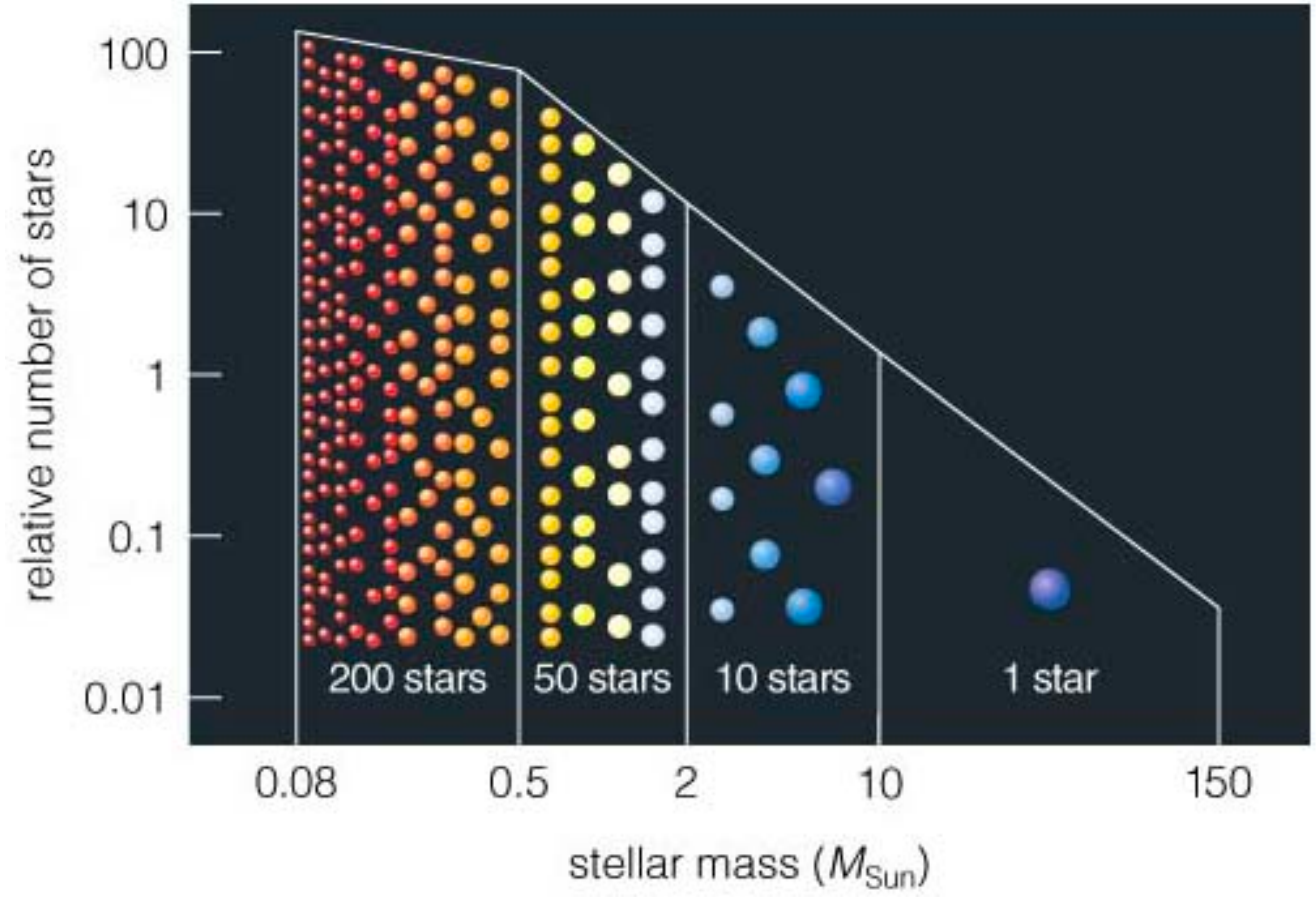
O star: $M = 60 \ M_\odot, \ L \approx 2 \times 10^4 \ L_{\odot,V}$

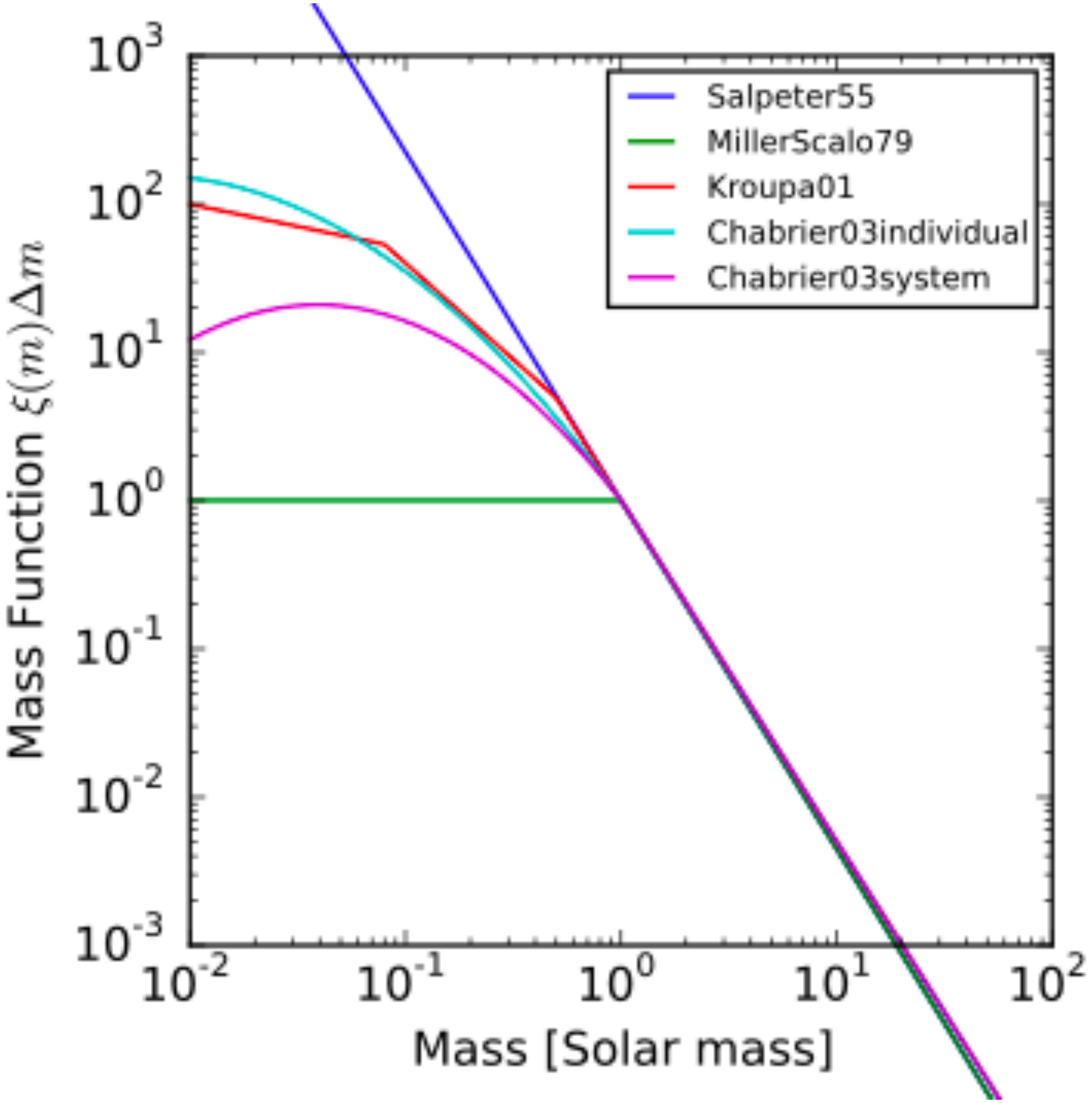
$$\langle M/L_V \rangle = 0.003 \ M_\odot/L_{\odot,V}$$

M star: $M = 0.1 \ M_\odot, \ L \approx 5 \times 10^{-5} \ L_{\odot,V}$

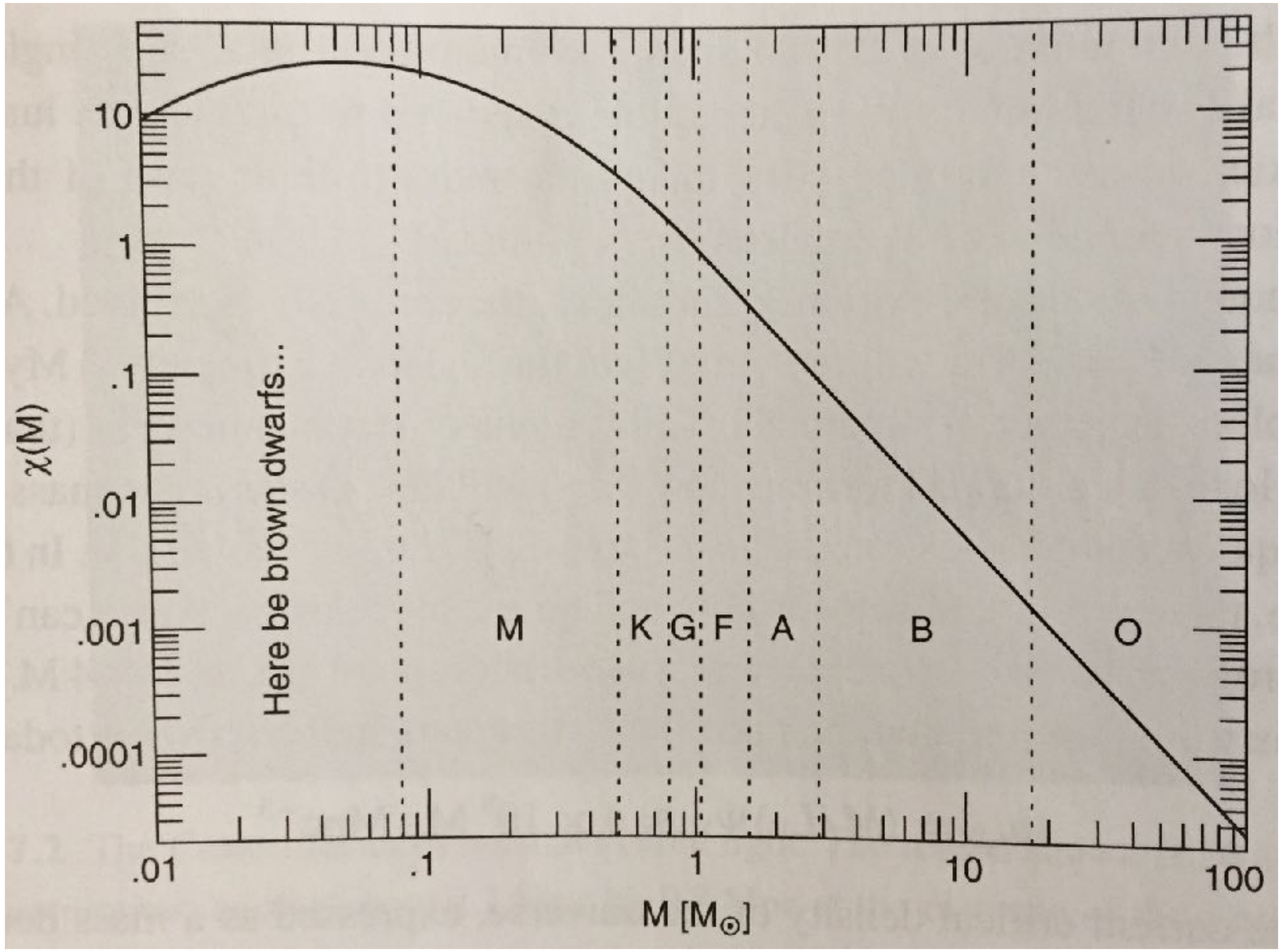
$$\langle M/L_V \rangle = 2000 \ M_\odot/L_{\odot,V}$$

Mass function of stars





Mass function of stars



$M < 1\ M_\odot$

$$\chi(M) \propto \frac{1}{M} \exp\left(-\frac{(\log M - \log M_c)^2}{2\sigma^2}\right)$$

$M > 1\ M_\odot$

$$\chi(M) \propto M^{-\beta}$$

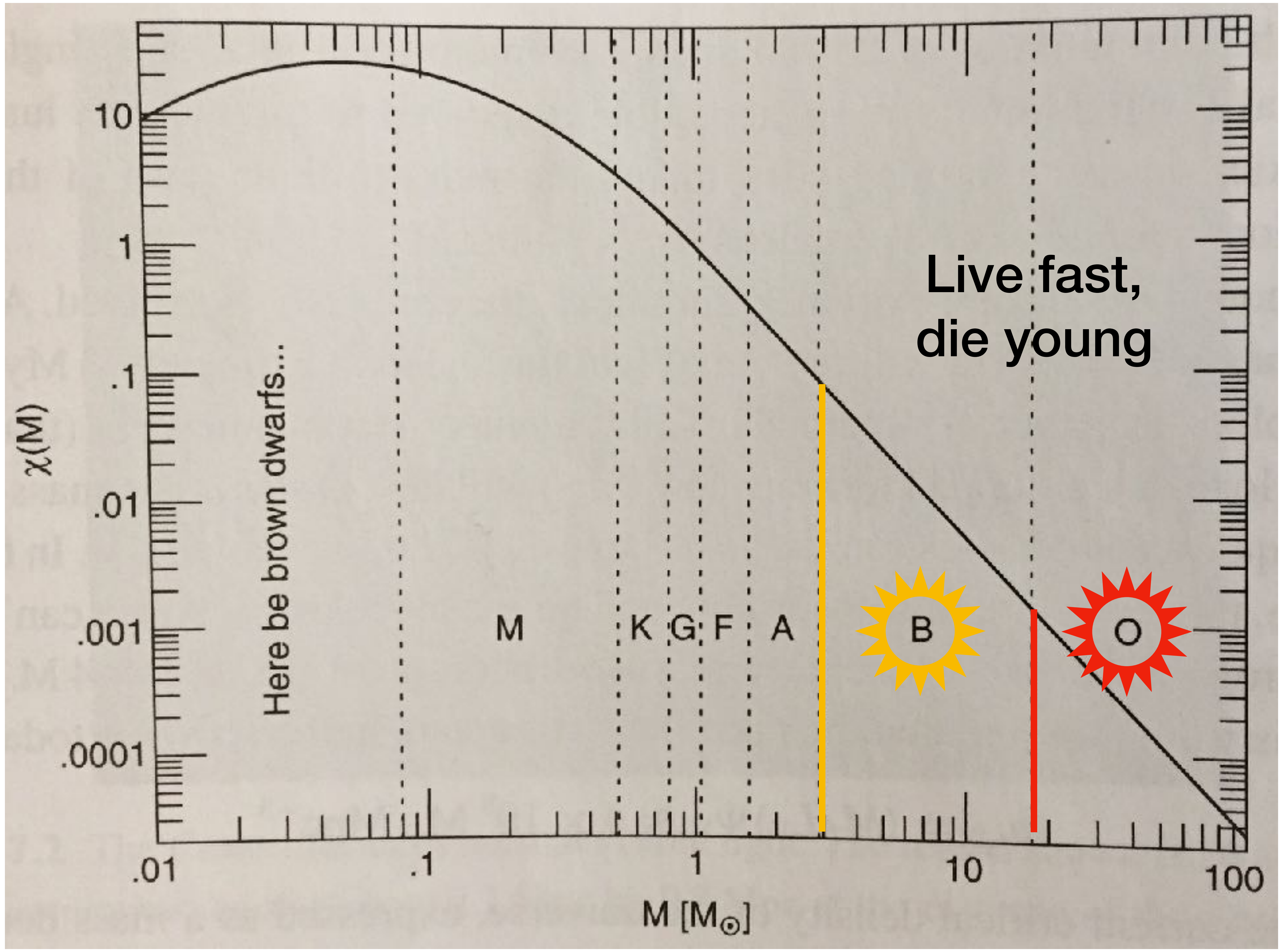
1 60 $M_\odot$ O star for every

10^5 0.2 $M_\odot$ M stars

but most light from O star

$\langle M/L_{\odot,V} \rangle \approx 0.3\ M_\odot/L_{\odot,V}$

Mass function of stars



quiescent (red and dead) galaxies

$$\langle M/L_{\odot,V} \rangle \approx 8\ M_\odot/L_{\odot,V}$$

mix of quiescent & SF galaxies

$$\rho_{*,0} = \langle M/L_{\odot,V} \rangle \Psi_V$$

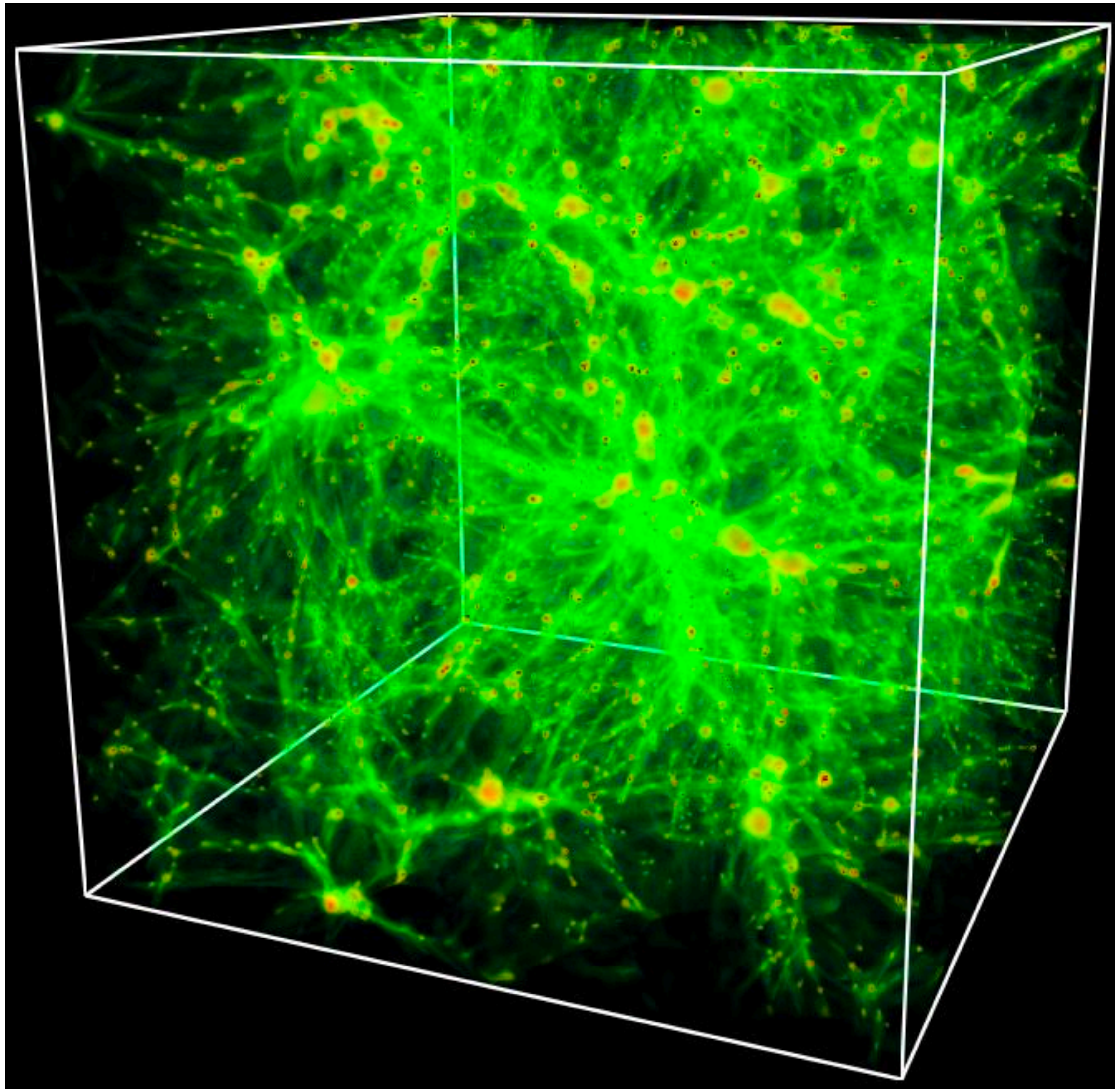
$$\approx 4 \times 10^8\ \mathrm{M}_\odot\ \mathrm{Mpc}^{-3}$$

$$\Omega_{*,0} = \frac{\rho_{*,0}}{\rho_{*,c}}$$

$$\approx \frac{4 \times 10^8\ \mathrm{M}_\odot\ \mathrm{Mpc}^{-3}}{1.28 \times 10^{11}\ \mathrm{M}_\odot\ \mathrm{Mpc}^{-3}}$$

$$\approx 0.003$$

Not all baryons are in stars, however

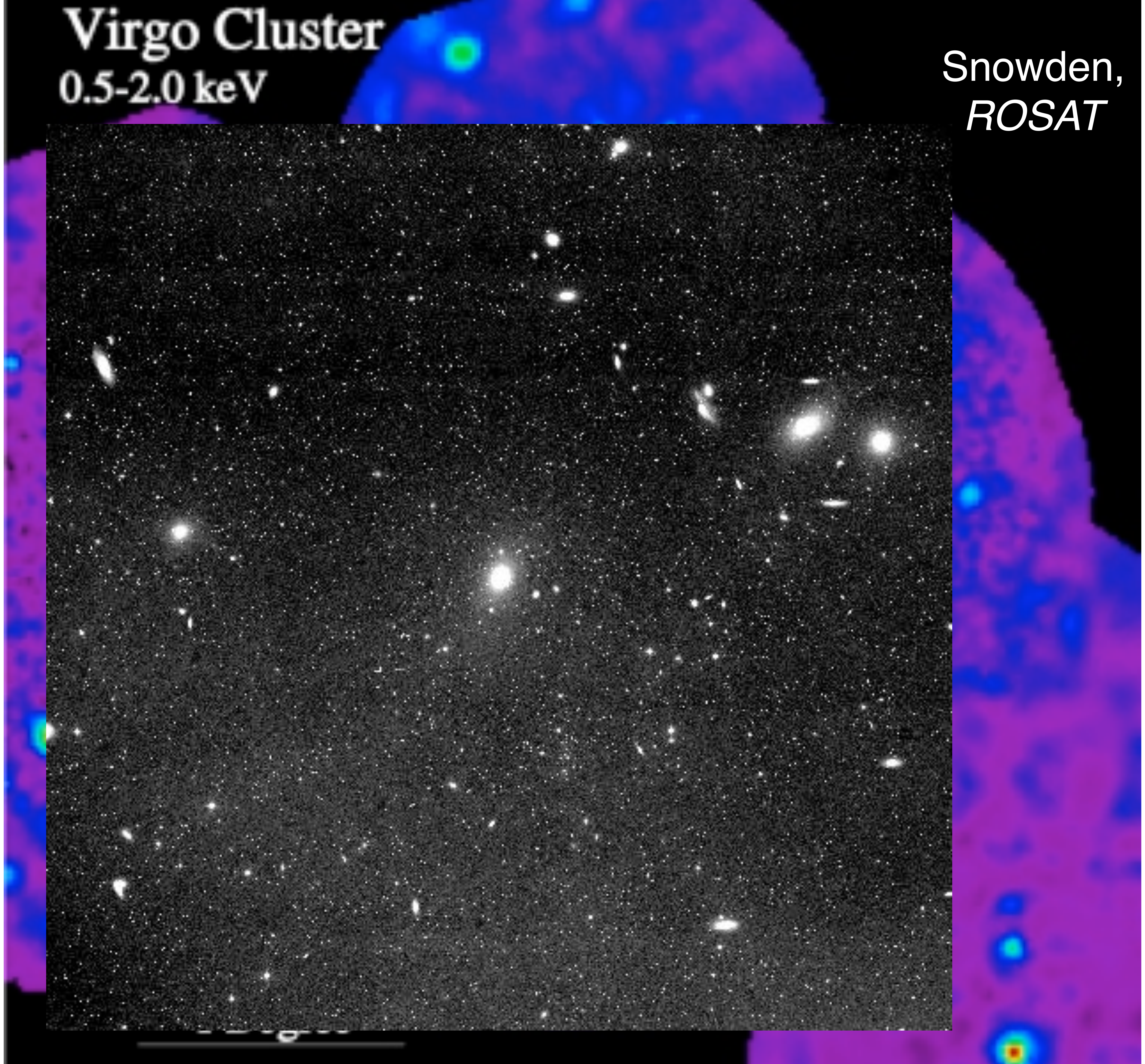


cosmological simulation showing the “warm-hot” gas in between galaxies in intergalactic space

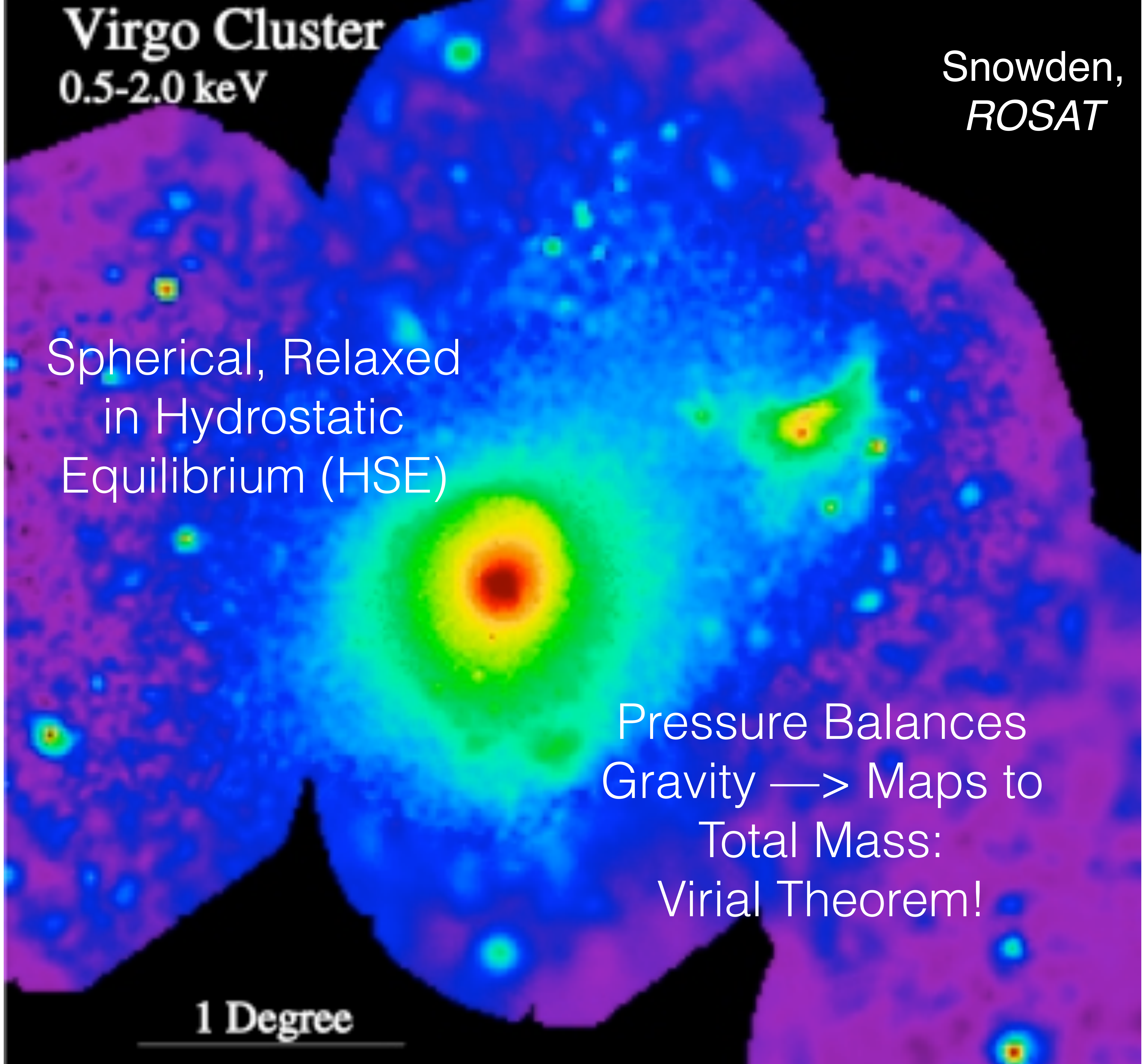
Galaxy

Group of Galaxies

Group of Galaxies



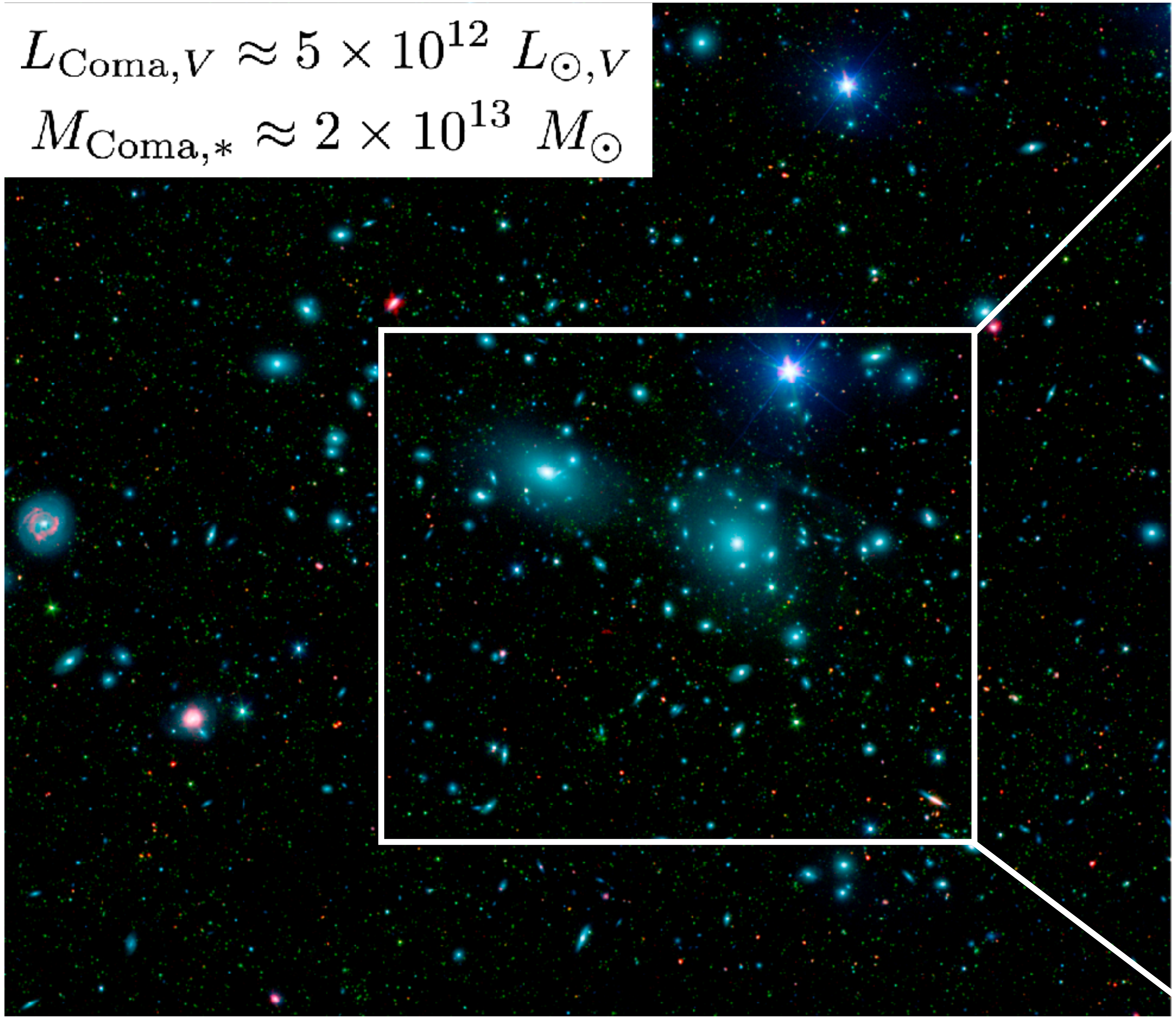
Snowden,
ROSAT


Virgo Cluster
0.5-2.0 keV
Snowden,
ROSAT
Spherical, Relaxed
in Hydrostatic
Equilibrium (HSE)
Pressure Balances
Gravity —> Maps to
Total Mass:
Virial Theorem!
1 Degree

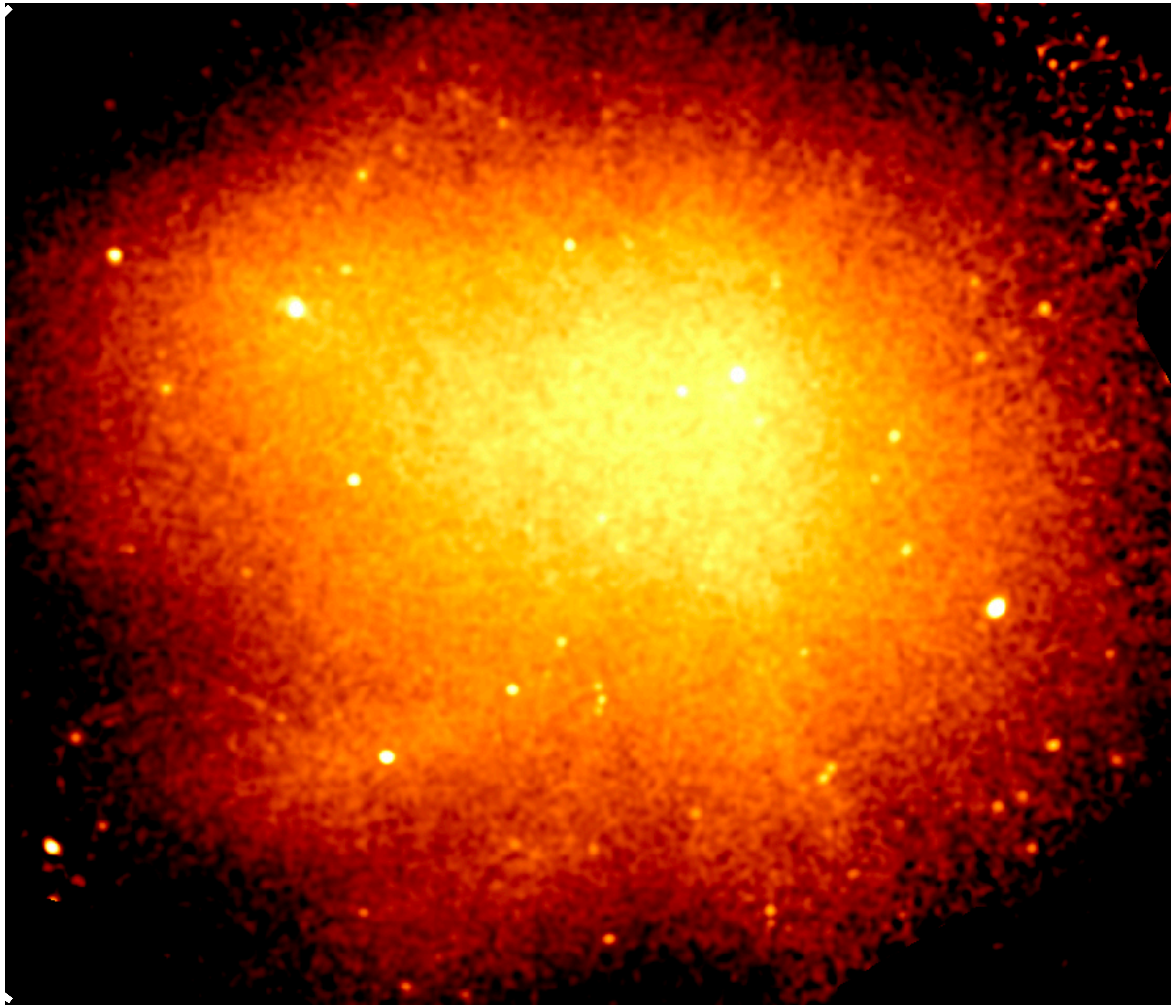
The Coma Cluster

$$L_{\mathrm{Coma},V} \approx 5 \times 10^{12}\ L_{\odot,V}$$

$$M_{\mathrm{Coma},*} \approx 2 \times 10^{13}\ M_{\odot}$$



$$M_{\mathrm{Coma,gas}} \approx 2 \times 10^{14}\ M_{\odot}$$



Baryonic Matter

$$\Omega_{*,0} \lesssim 0.005$$

$$M_{\mathrm{gas},0} \approx 10 \times M_{*,0}$$

early universe measurements

$$\Omega_{\mathrm{bary},0} = 0.048 \pm 0.003$$

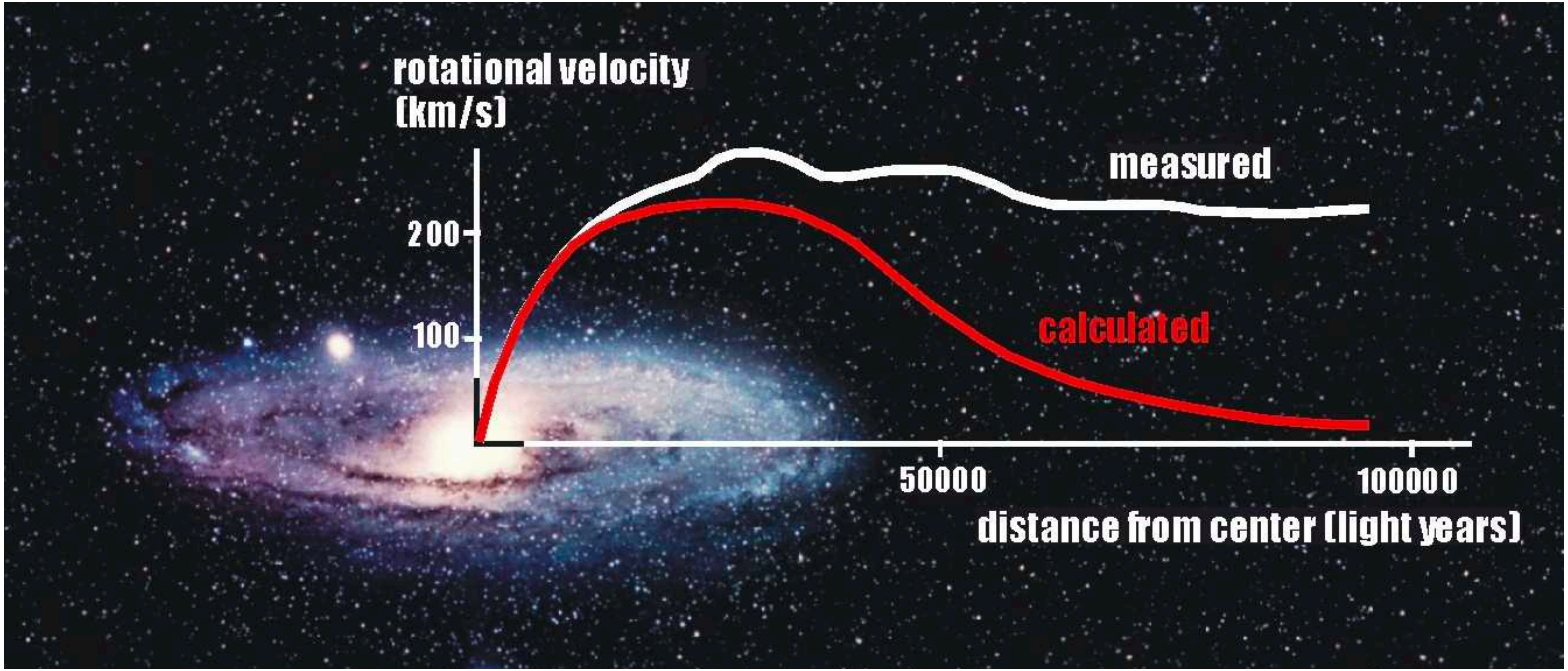
$$\Omega_{m,0} = 0.31$$

baryonic matter only 15%

By the time of the Big Bang and thereafter, normal matter is the subdominant form of matter in the universe, with some other form of matter (non-baryonic dark matter) making up the majority of non-relativistic matter in the universe

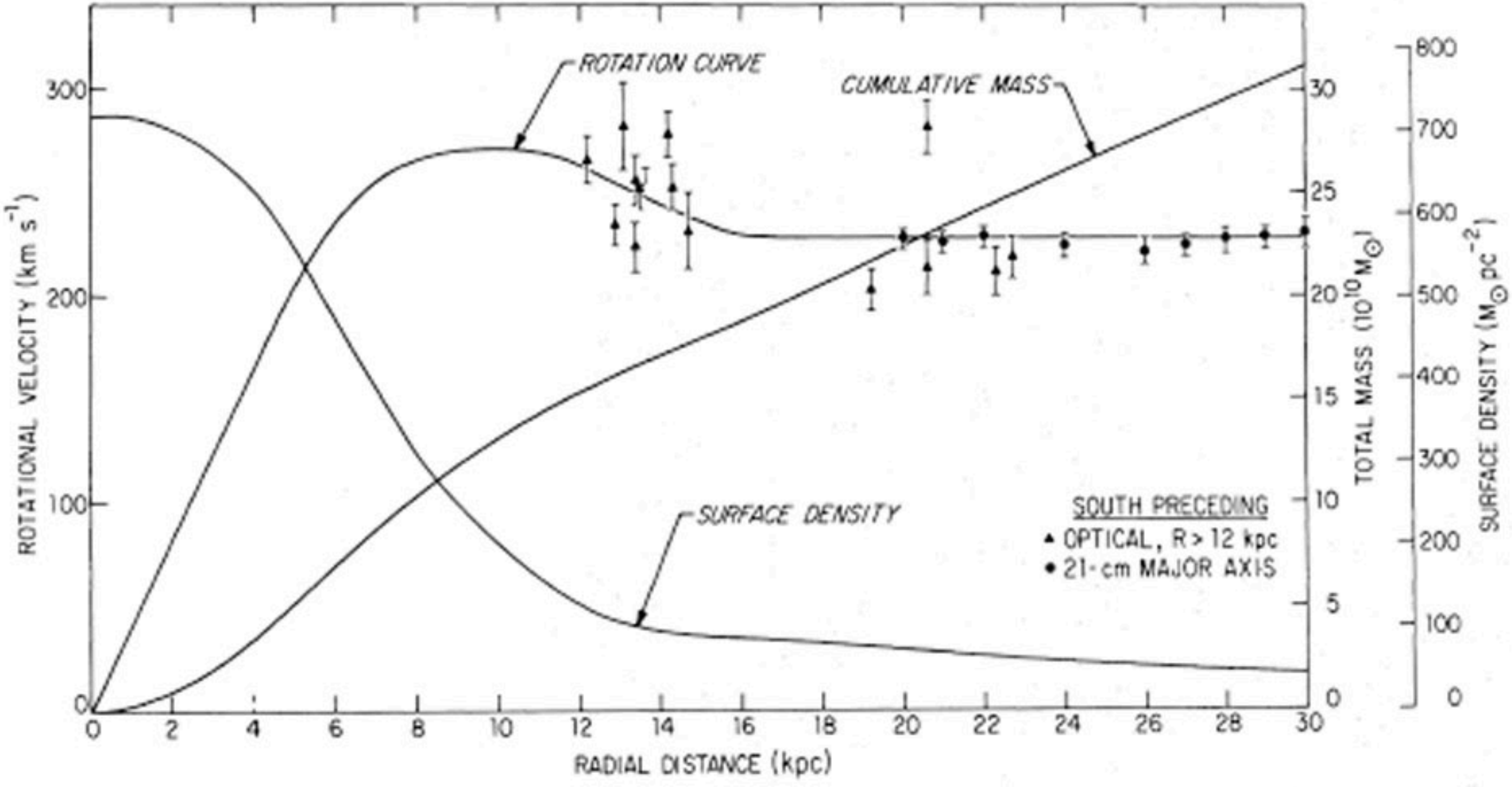
Could be primordial black holes that were made before this time (i.e., not from stars).

Dark Matter in Galaxies

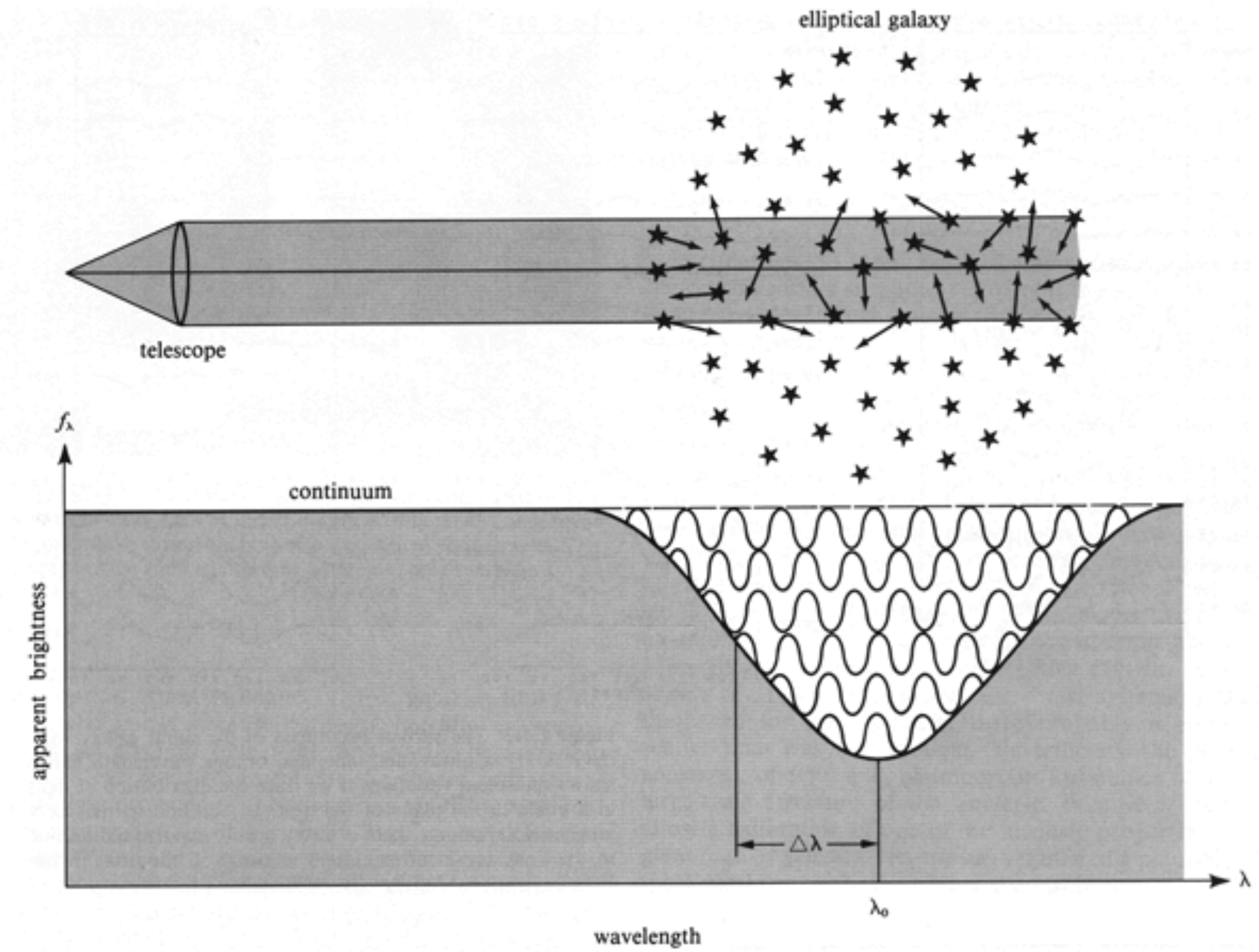


$$M(R) = \frac{v^2 R}{G} = 1.05 \times 10^{11}\ \mathrm{M}_\odot \left(\frac{v}{235\ \mathrm{km\ s}^{-1}} \right)^2 \left(\frac{R}{8.2\ \mathrm{kpc}} \right)$$

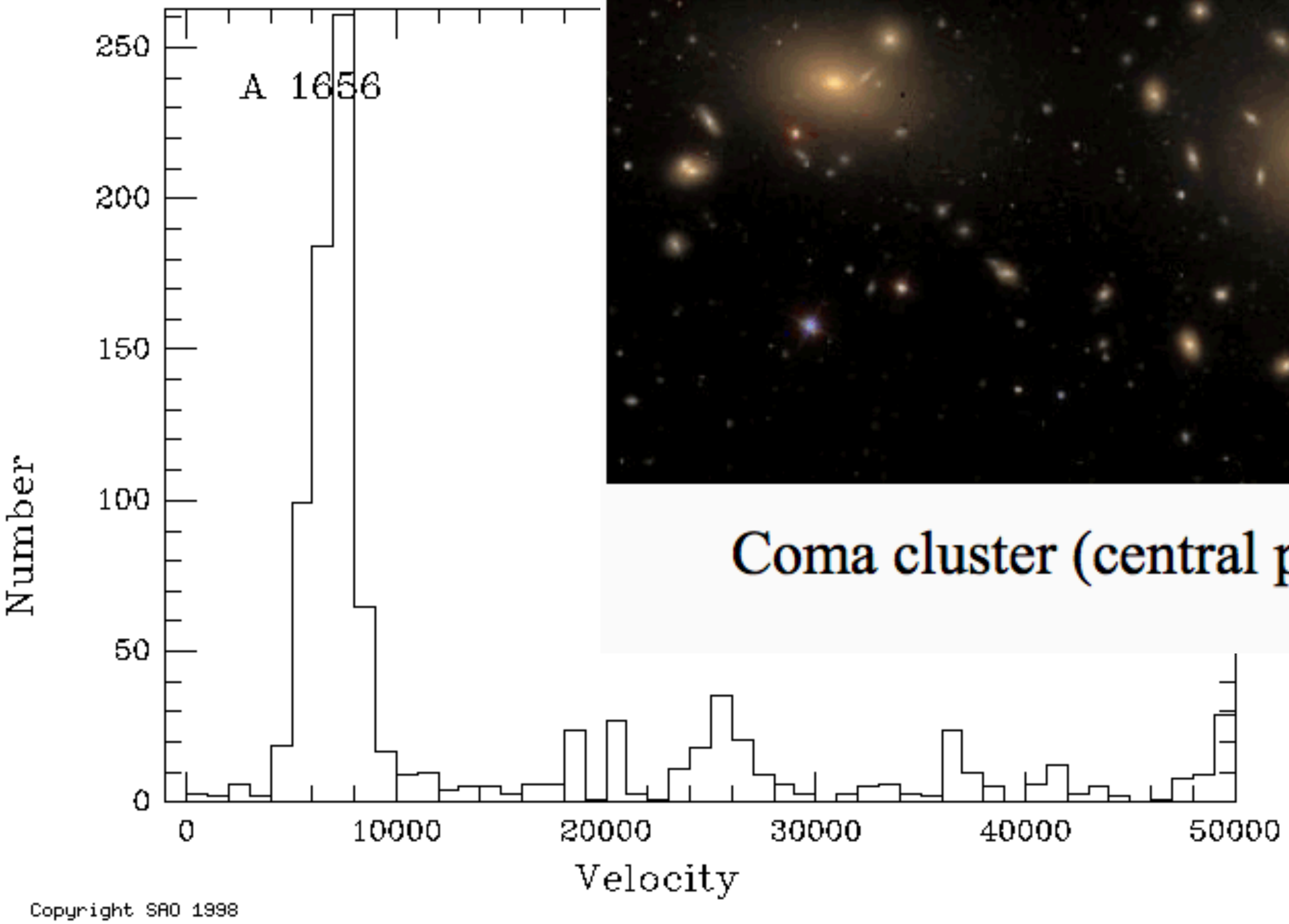
Dark Matter in Galaxies



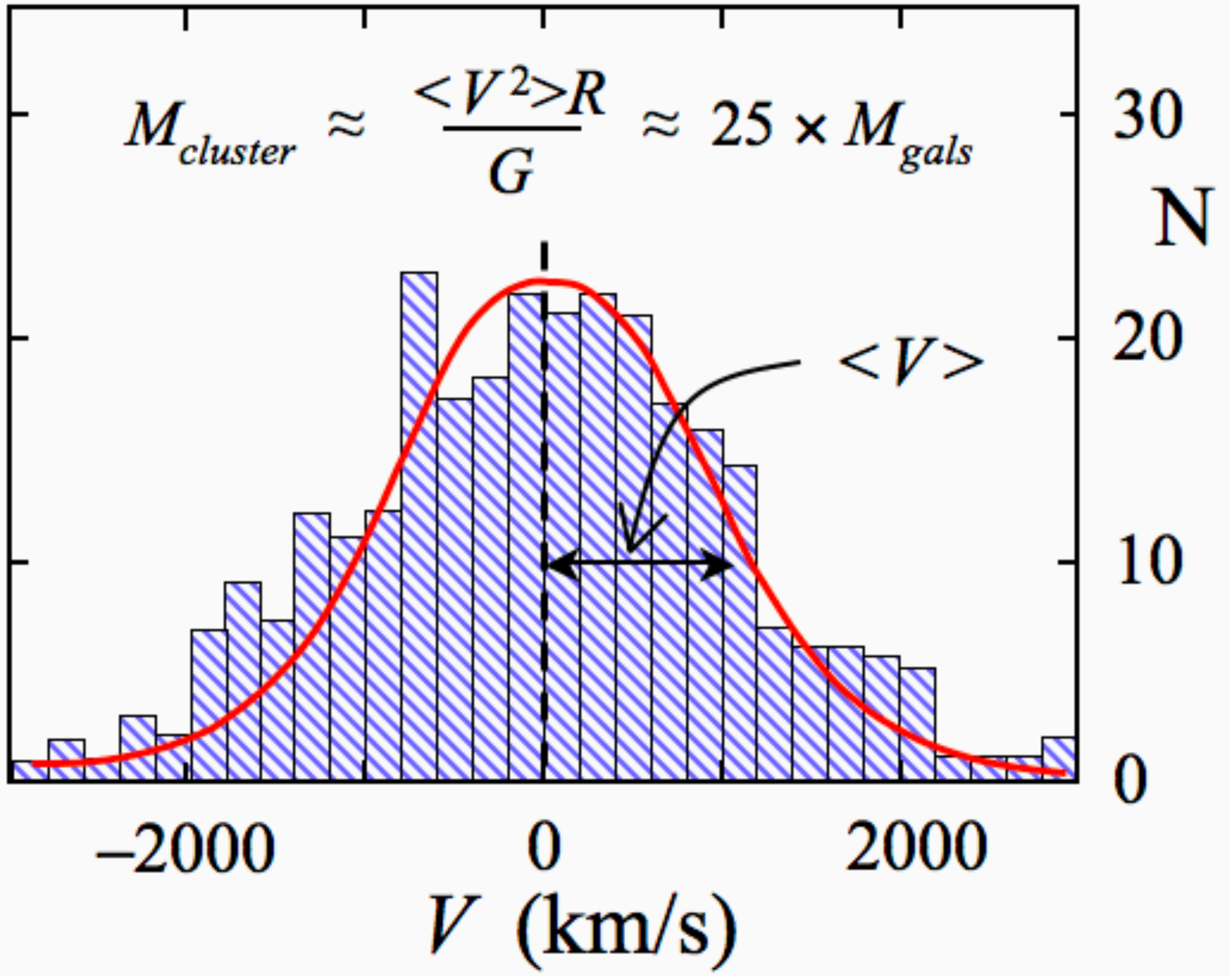
Dark Matter in Clusters



Dark Matter in Clusters



Coma cluster (central part)



Total mass from the hot gas

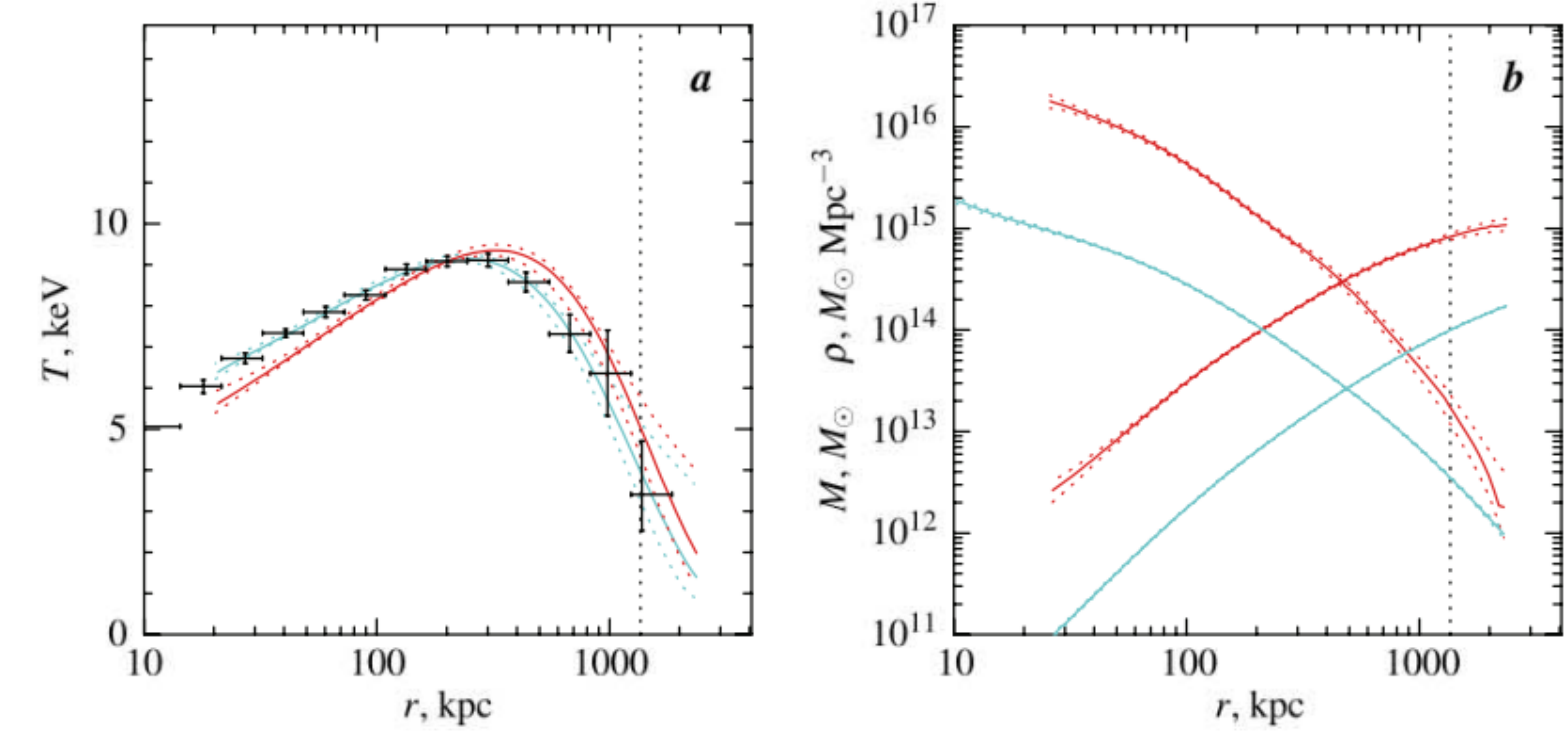
$$\frac{dP_{\rm gas}}{dr} = -\frac{GM(r)\rho_{\rm gas}(r)}{r^2}$$

$$P_{\rm gas} = \frac{\rho_{\rm gas} kT_{\rm gas}}{\mu}$$

$$M(<r) = \frac{kT_{\rm gas}(r)r}{G\mu}\left[-\frac{d\ln\rho_{\rm gas}}{d\ln r} - \frac{d\ln T_{\rm gas}}{d\ln r}\right]$$

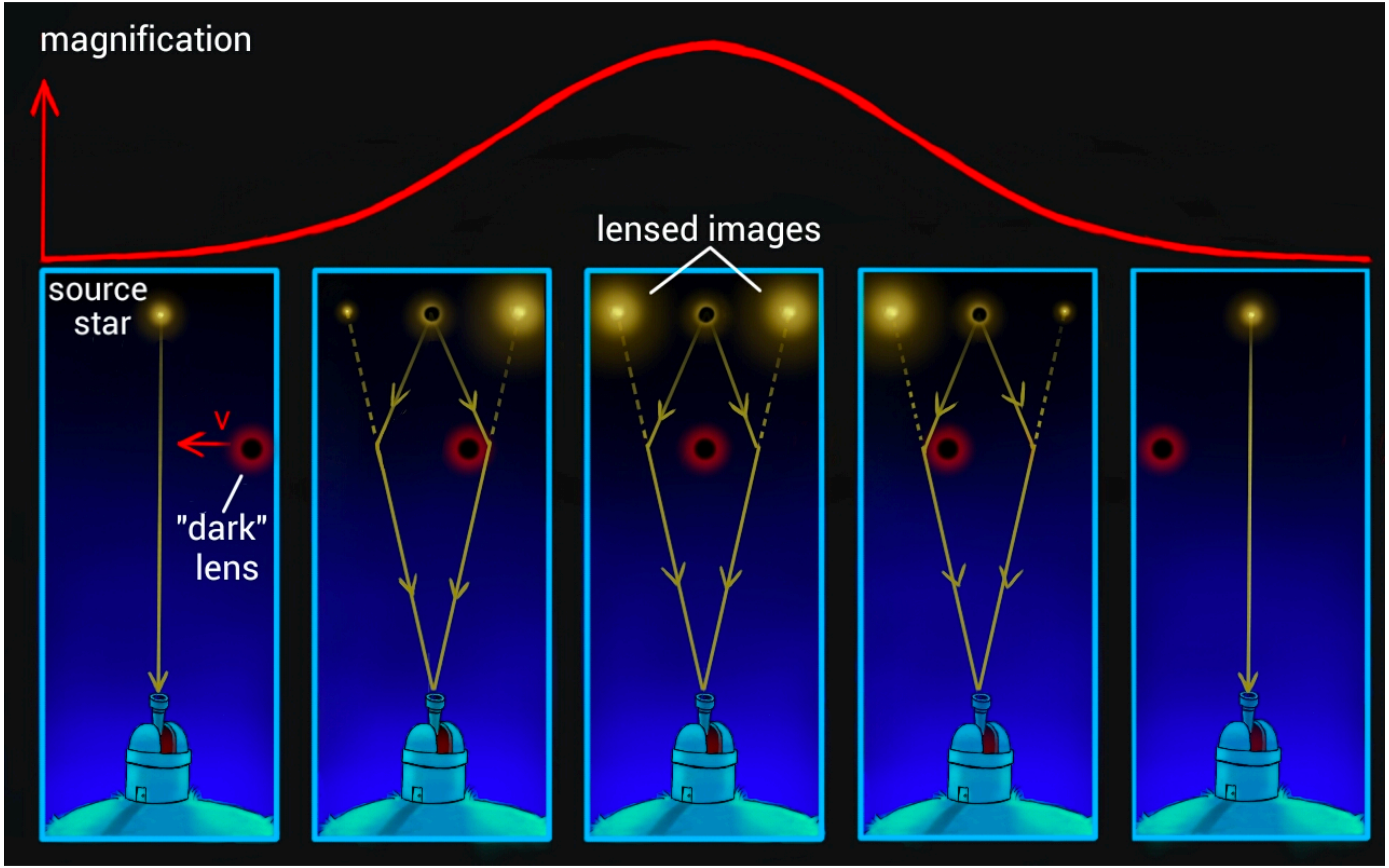
Total mass of clusters alone yield —> $\Omega_{\rm clus,0} \approx 0.2$ (lower limit on $\Omega_{m,0}$)

Total mass from the hot gas

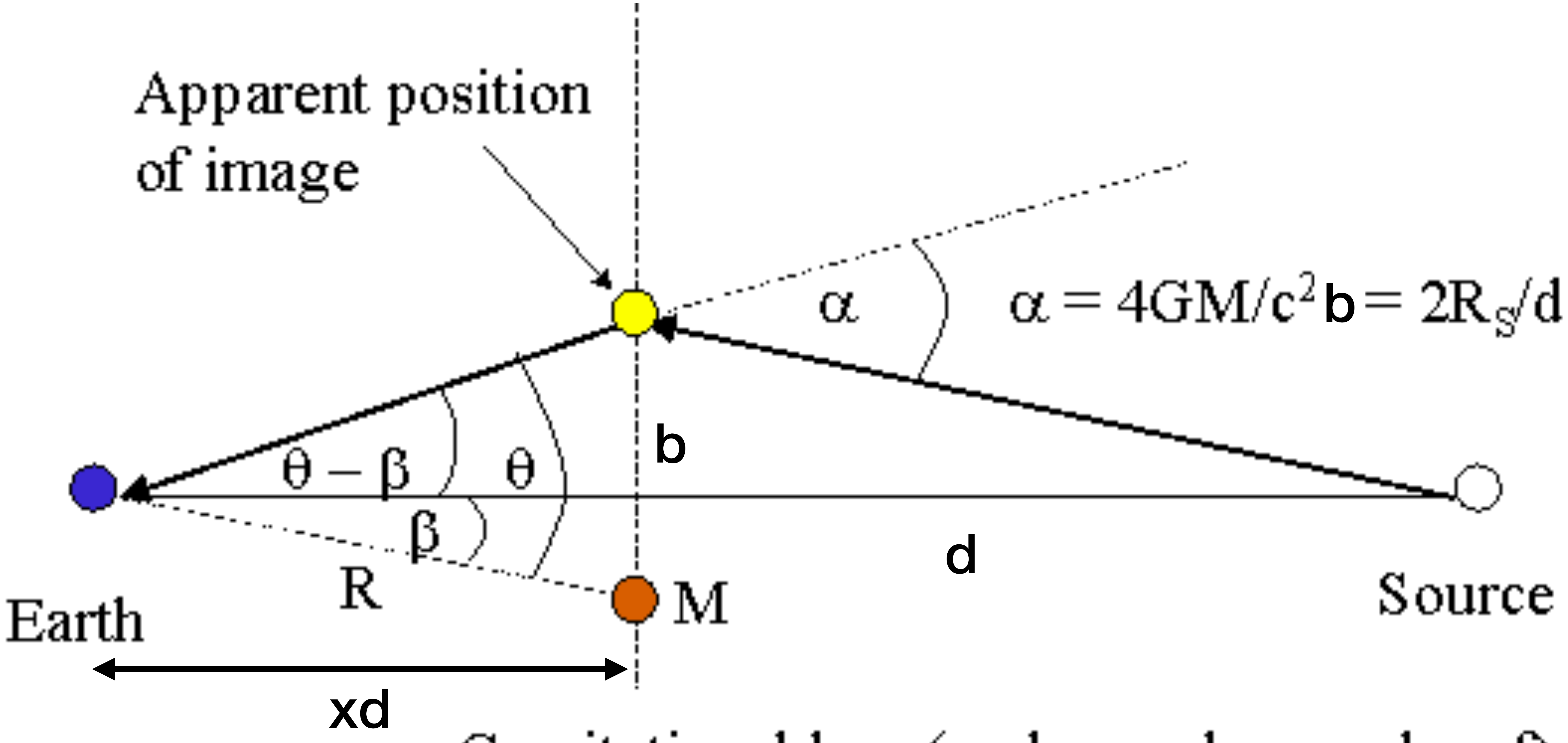


Detecting MACHOs via gravitational lensing

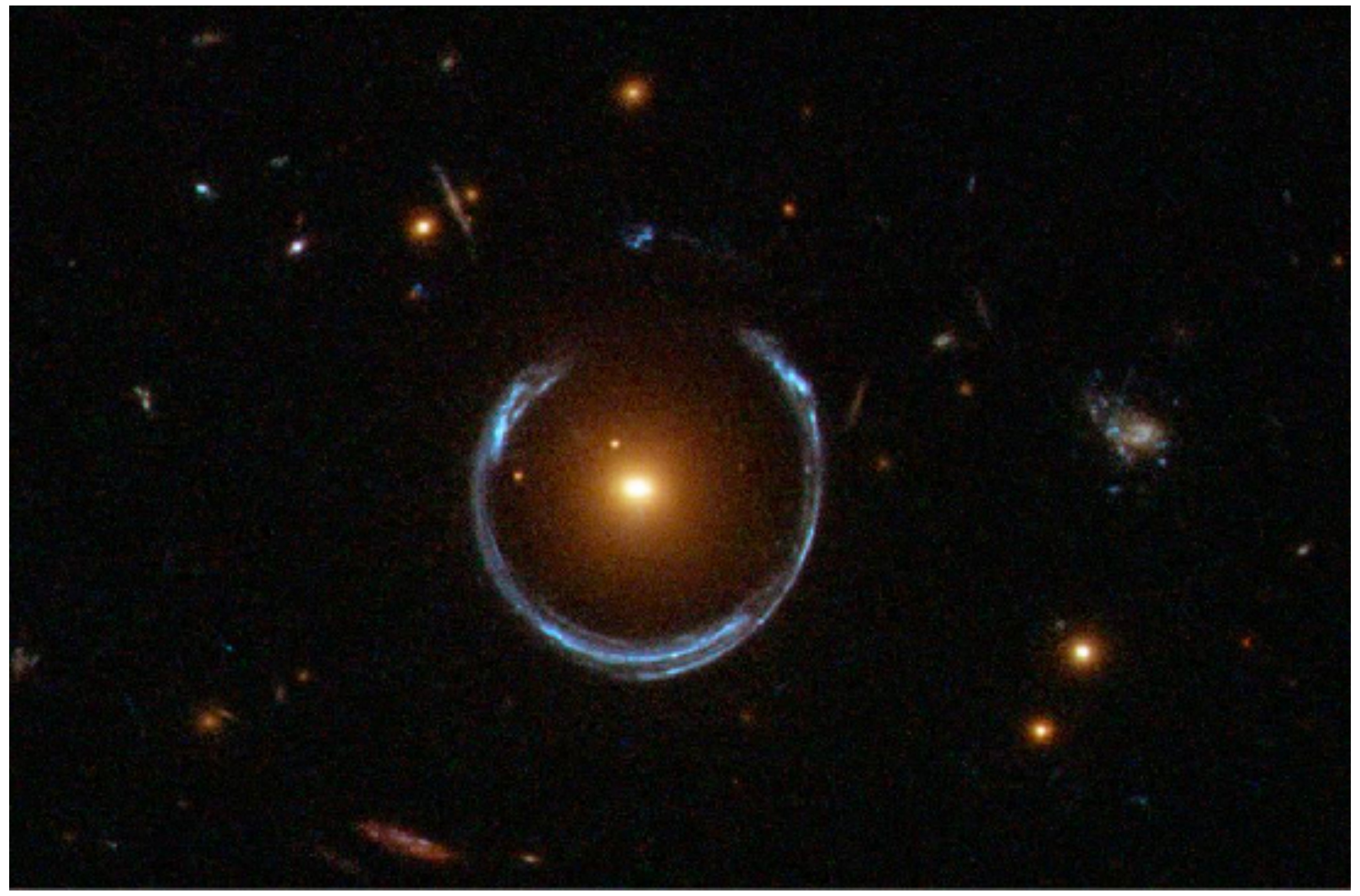
MAssive
Compact
Halo
Object



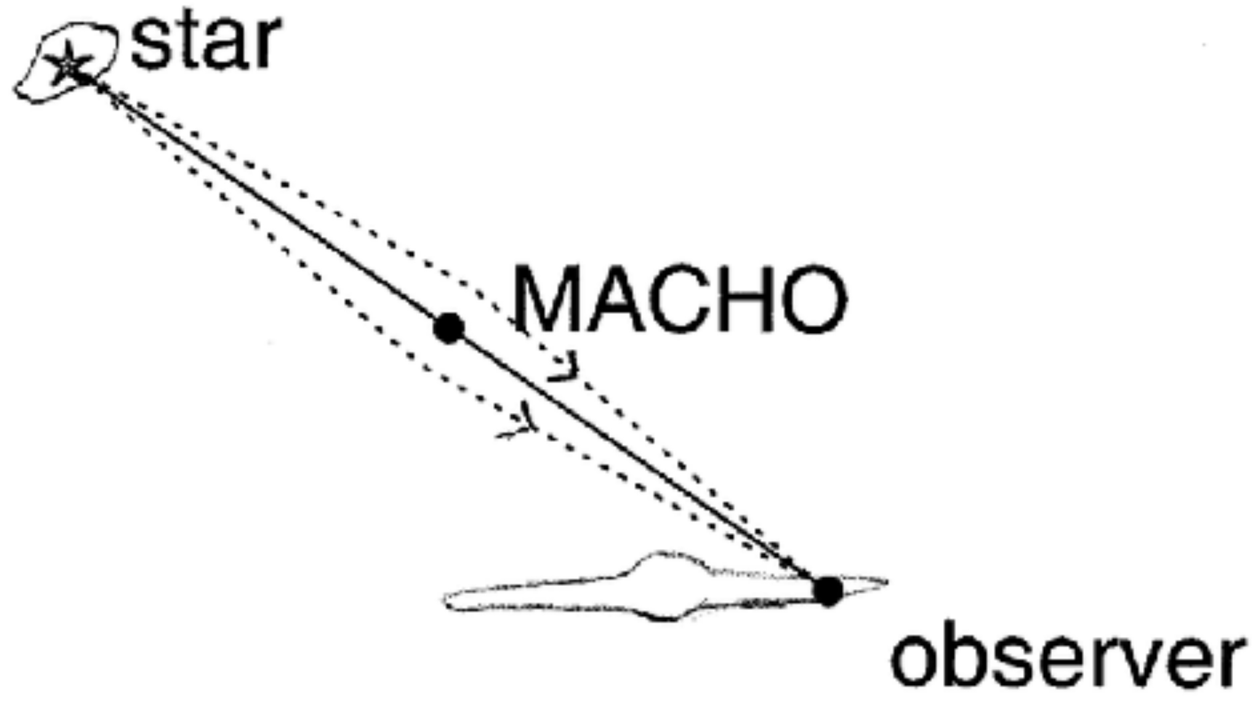
Detecting MACHOs via gravitational lensing



Gravitational lens (perhaps a brown dwarf)



$$\theta_E = \left(\frac{4GM}{c^2 d}\frac{1-x}{x}\right)^{1/2}$$



R = distance to lens (of mass M)

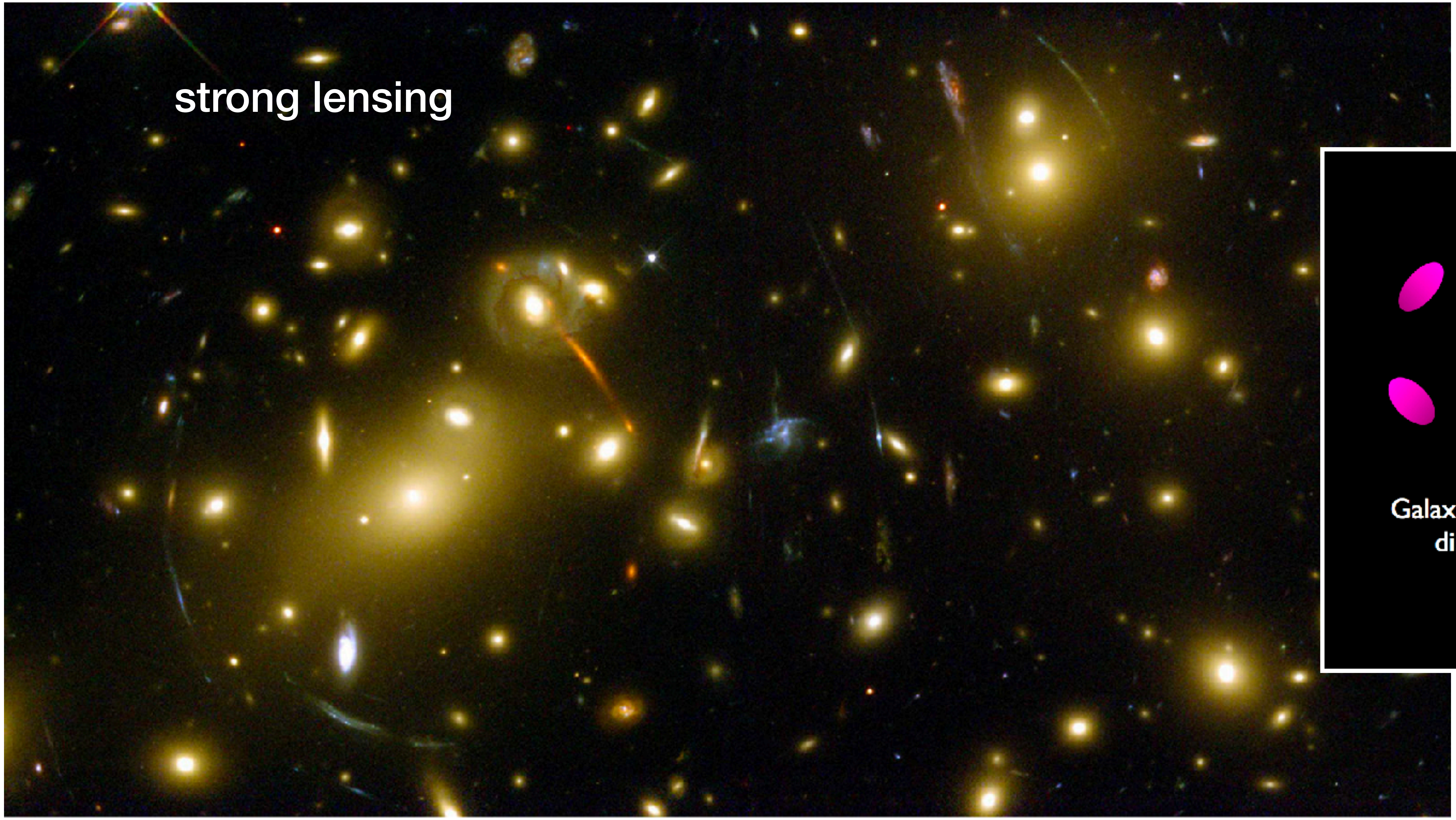
d = distance to source

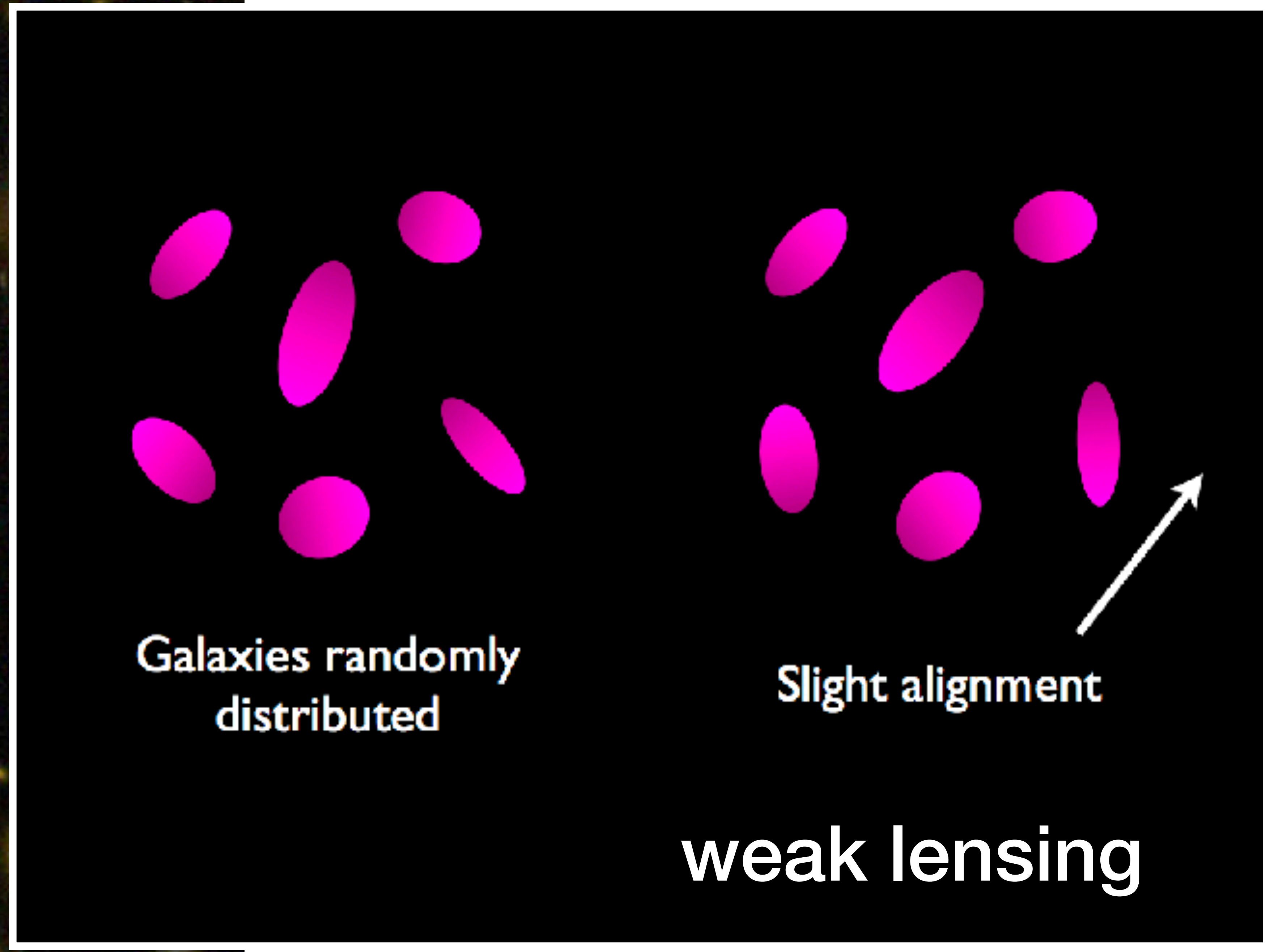
b = distance between lens and image

R_S= Schwarzchild radius of lens

$$\theta_E \approx 4\times 10^{-4}\ \text{arcsec}\ \times \left(\frac{M}{1\ M_\odot}\right)^{1/2}\left(\frac{d}{50\ \text{kpc}}\right)^{-1/2}$$

Gravitational lensing by galaxy clusters





What could (non-baryonic) dark matter be?

cosmic neutrinos?

in the Standard Model, neutrinos are massless (but we now know that's not the case)

their number density is set by early universe calculations,
so knowing their mass yields their density parameter

constraints on their mass:

$$0.019 \text{ eV} < m_\nu c^2 < 0.1 \text{ eV}$$

lead to constraints on the density parameter:

$$0.0013 < \Omega_{\nu,0} < 0.007$$

DEPARTMENT OF PHYSICS & ASTRONOMY

COLLOQUIUM

HET FACULTY CANDIDATE

IAN SHOEMAKER

UNIVERSITY OF SOUTH DAKOTA

NEUTRINOS, DARK MATTER, AND THE QUEST FOR NEW PHYSICS

The capstone discovery of the Higgs Boson marked the confirmation of the final piece of the well-tested Standard Model of particle physics. Despite the incredible successes of the Standard Model, it cannot be the complete description of nature. Chief among the observational facts requiring new physics are the existence of neutrino masses and non-luminous "Dark Matter" which dominates the Universe's matter budget. In this talk I'll discuss theoretical extensions of the Standard Model involving dark matter and neutrinos, along with the broad experimental program currently underway that will help uncover which the new physics framework nature has chosen.

Thursday February 22, 2018
4:00 pm
Room 102 JFB

Refreshments at 3:30 pm in JFB 219.

DEPARTMENT OF PHYSICS & ASTRONOMY

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YUE ZHOU

UNIVERSITY OF MICHIGAN

DARK MATTER BEYOND WEAKLY INTERACTING MASSIVE PARTICLES

Dark matter (DM) comprises approximately 27% of the energy in the observable universe. Its properties, such as its mass and interactions, remain largely unknown. Unveiling the properties of DM is one of the most important tasks in high energy physics. For the past few years, motivated by possible new physics at the electroweak scale, many DM experiments have looked for DM with mass at O(100) GeV. This is not the only possibility, however. Large chunks of parameter space supported by other well-motivated models remain to be carefully studied. Exploring these regimes requires creative ideas and advanced technologies. I will first talk about a novel proposal using superconductors as the target material for DM direct detection. This setup has the potential to lower the direct detection mass threshold from a few GeV to keV, consequently probing the warm dark matter scenario. Then I will present a recent proposal utilizing the Gravitational Wave (GW) experiments, i.e., LIGO and LISA, to search for ultra-light dark photon dark matter. We show these GW experiments can go well beyond existing constraints and probe large regions of unexplored parameter space. Both proposals are under serious investigation by experimental groups and likely to be carried out in the near future.

Thursday March 1, 2018
4:00 pm
Room 102 JFB

Refreshments at 3:30 pm in JFB 219.

Non-baryonic dark matter candidates

WIMPs

Weakly Interacting Massive Particles
(supersymmetric extension of the Standard Model)

Axions

(hypothetical particle that explains why quantum chromodynamics does not “break CP symmetry”)

Sterile Neutrinos

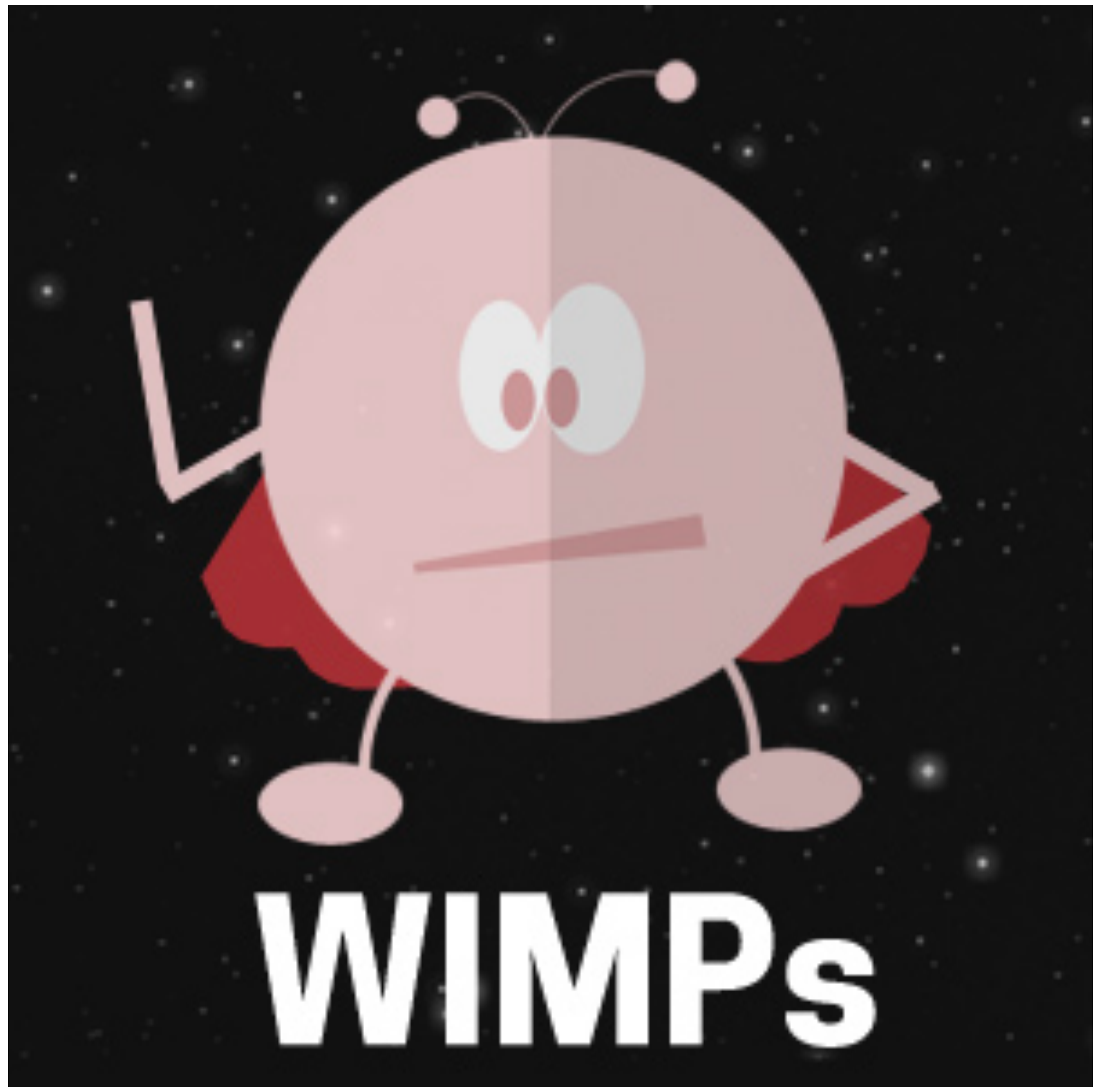
(right handed partner to known neutrinos, but doesn’t experience weak force interactions)

WIMPs

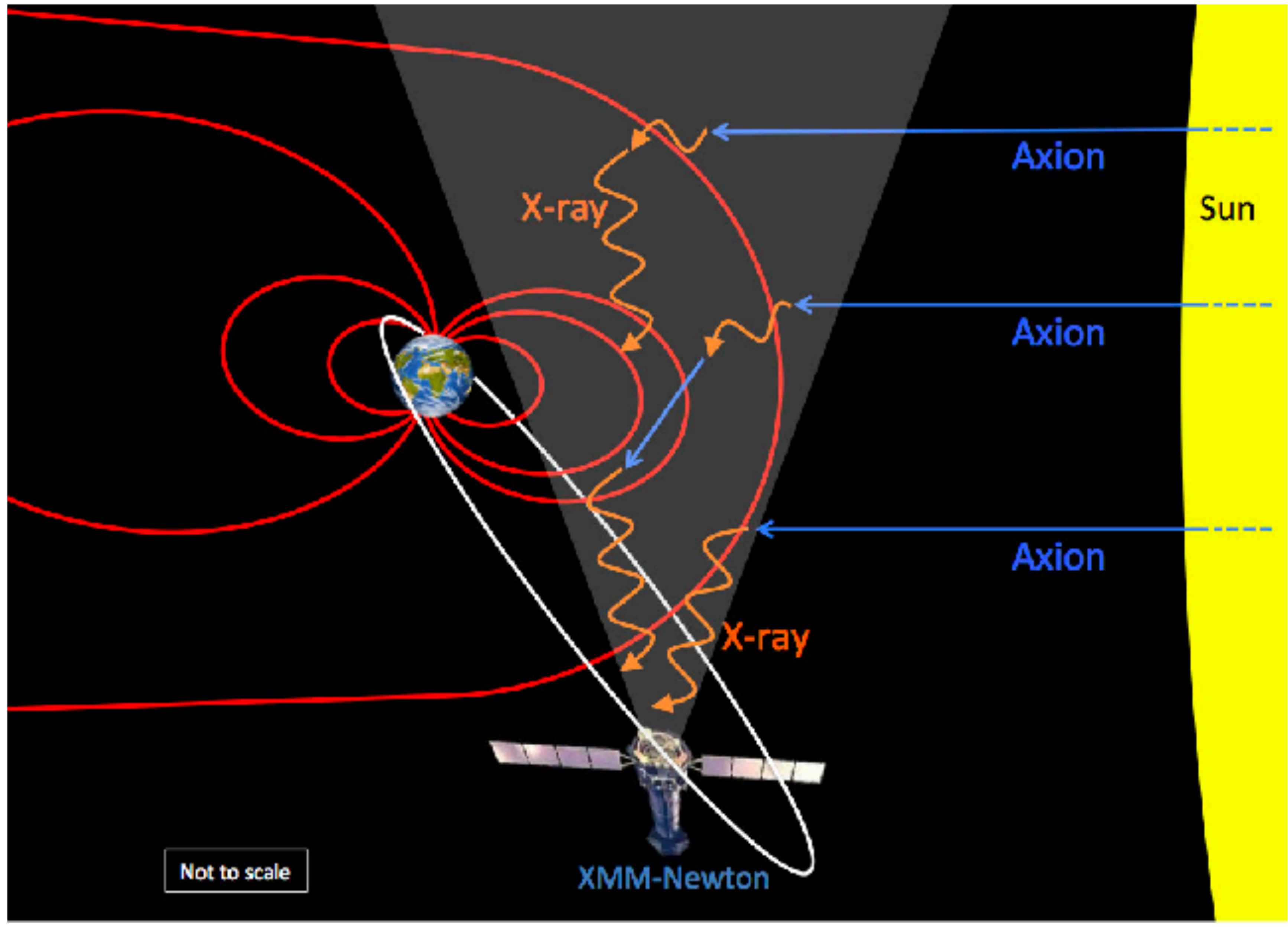
very loosely defined (any new particle that's relatively massive and interacts via gravity [and potentially other sources])

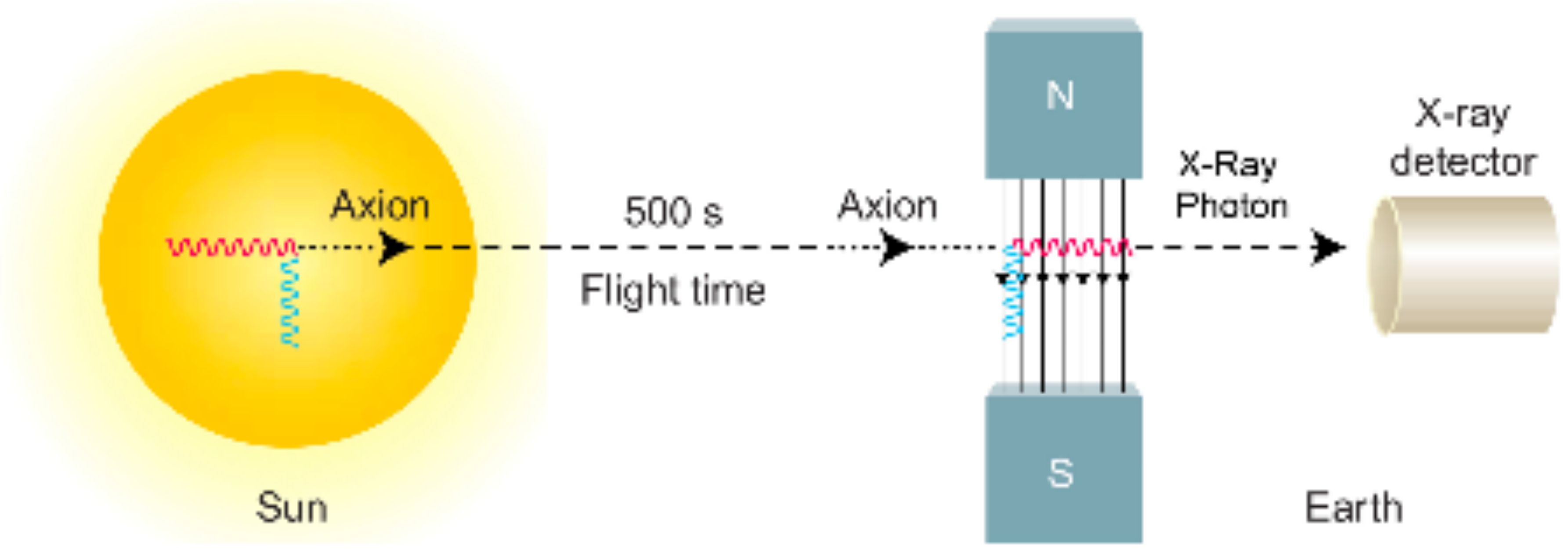
supersymmetric extensions of the SM (positing more massive versions of all known particles) naturally lead to WIMP production in the Big Bang —> called the "WIMP miracle" (direct detection searches and the LHC have failed to find WIMPs at these "miraculous" masses)

their self-annihilation (into gamma ray photons) could be detected in dark matter concentrations, such as the centers of galaxies and clusters of galaxies
(no definitive observations — without other reasonable explanations — have been made)



Axions





in QCD, strong interactions permit violations of charge conjugation (that if you swap the charge signs of particles and anti-particles, the laws of physics remain unchanged) and parity (no "handedness" in interactions)

—> would lead to an electric dipole moment for the neutron, which has been measured to be consistent with zero (with an upper limit making it very small)

—>—> this requires a term, which in SM theory could be any number b/t 0 and 2pi, to be very close to 0, and by "naturalness" arguments this is a "problem"

—>—>—> can be solved if there's a new particle (the axion) that could also serve as a dark matter particle

original version of the axion has been ruled out by experiment

current dark matter axion candidates are variations on this idea, but not as well motivated by theory

can be converted into photons in a strong magnetic field and detected that way

Sterile Neutrinos

"sterile" because they don't interact via the weak force like SM neutrinos

right-handed chirality (spin vector relative to momentum)
SM particles have left and right varieties, SM neutrinos are left-handed only

can have any mass (1 eV to 10^{15} GeV)

their decay would produce 2 photons (each with half the energy of the neutrino, which for dark matter would have to be non-relativistic so $E=mc^2$)

detection (and non-detections) at X-ray (keV) energies